

Test of Einstein Equivalence Principle for 0-Spin and Half-Integer-Spin Atoms: Search for Spin-Gravity Coupling Effects

M. G. Tarallo,^{*} T. Mazzoni, N. Poli, D. V. Sutyryn, X. Zhang,[†] and G. M. Tino[‡]

*Dipartimento di Fisica e Astronomia and LENS—Università di Firenze, INFN—Sezione di Firenze,
Via Sansone 1, 50019 Sesto Fiorentino, Italy*

(Received 24 February 2014; published 8 July 2014)

We report on a conceptually new test of the equivalence principle performed by measuring the acceleration in Earth's gravity field of two isotopes of strontium atoms, namely, the bosonic ^{88}Sr isotope which has no spin versus the fermionic ^{87}Sr isotope which has a half-integer spin. The effect of gravity on the two atomic species has been probed by means of a precision differential measurement of the Bloch frequency for the two atomic matter waves in a vertical optical lattice. We obtain the values $\eta = (0.2 \pm 1.6) \times 10^{-7}$ for the Eötvös parameter and $k = (0.5 \pm 1.1) \times 10^{-7}$ for the coupling between nuclear spin and gravity. This is the first reported experimental test of the equivalence principle for bosonic and fermionic particles and opens a new way to the search for the predicted spin-gravity coupling effects.

DOI: 10.1103/PhysRevLett.113.023005

PACS numbers: 37.25.+k, 03.75.Dg, 04.80.Cc, 37.10.Jk

The Einstein equivalence principle (EP) is at the heart of general relativity, the present theory of gravity [1]. In its so-called *weak* form, corresponding to the universality of free fall, it goes back to Galileo Galilei's idea that the motion of a mass in a gravitational field is independent of its structure and composition. Violations of the EP are expected in attempts to unify general relativity with the other fundamental interactions and in theoretical models for dark energy in cosmology [2,3] as well as in extended theories of gravity [4].

The most stringent experimental limits for the EP to date come from two methods: the study of the motion of moons and planets and the use of torsion balances [5]. In recent years, experiments based on atom interferometry [6,7] compared the fall in Earth's gravitational field of two Rb isotopes [8,9] and Rb versus K [10] reaching a relative precision of about 10^{-7} . Tests of EP were carried out in which the measurement of Earth's gravity acceleration with an atom interferometer was compared with the value provided by a classical gravimeter [11,12]. A much higher precision will be achieved in future experiments with atom interferometers that are planned on the ground [13] and in space [14,15]. The possibility of tests with atom interferometry for matter versus antimatter was also investigated [16,17]. The interest of using atoms is indeed not only to improve the limits reached by classical tests with macroscopic bodies, but mostly in the possibility to perform qualitatively new tests with "test masses" having well-defined properties, e.g., in terms of spin, bosonic or fermionic nature, and proton-to-neutron ratio.

Possible spin-gravity coupling, torsion of space-time, and EP violations have been the subject of extensive theoretical investigation (see, for example, Refs. [18–24]). Experimental tests were performed based on macroscopic test masses [24,25], atomic magnetometers [26,27], and

atomic clocks [28]. In Ref. [8], a differential free fall measurement of atoms in two different hyperfine states was also performed. Possible differences in gravitational interaction for bosonic and fermionic particles were also discussed [29,30] and efforts towards experimental tests with different atoms are under way [30,31].

In this Letter we report on an experimental comparison of the gravitational interaction for a bosonic isotope of strontium (^{88}Sr) which has zero total spin with that of a fermionic isotope (^{87}Sr) which has a half-integer spin. Sr in the ground state has a 1S_0 electronic configuration and the total spin corresponds to the nuclear spin I ($I_{87} = 9/2$). Gravity acceleration was measured by means of a genuine quantum effect, namely, the coherent delocalization of matter waves in an optical lattice. To compare gravity acceleration for the two Sr isotopes, we confined atomic wave packets in a vertical off-resonant laser standing wave and induced a dynamical delocalization by amplitude modulation (AM) of the lattice potential [12,32,33] at a frequency corresponding to a multiple ℓ of the Bloch frequency $\nu_B = F_g \lambda_L / 2h$, where h is the Planck constant, λ_L is the wavelength of the optical lattice laser (Fig. 1), and F_g is the gravitational force on the atomic wave packet.

In order to account for anomalous acceleration and spin-dependent gravitational mass, the gravitational potential can be expressed as

$$V_{g,A}(z) = (1 + \beta_A + kS_z)m_A g z, \quad (1)$$

where m_A is the rest mass of the atom, β_A is the anomalous acceleration generated by a nonzero difference between gravitational and inertial mass due to a coupling with a field with nonmetric interaction with gravity [17,34], k is a model-dependent spin-gravity coupling strength, and S_z is the projection of the atomic spin along gravity direction. k

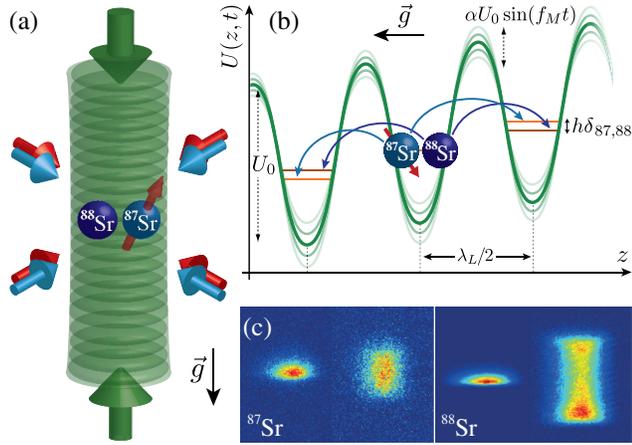


FIG. 1 (color online). Experimental configuration to test the equivalence principle with Sr atoms. (a) The two isotopes are alternately laser cooled and trapped in a vertical optical lattice. (b) Intraband coherent delocalization of atomic wave packets is induced by means of amplitude modulation of the optical lattice potential: the difference between the resonant modulation frequencies of the two atomic species $\delta = \nu_{B,88} - \nu_{B,87}$ depends only on their mass ratio and the EP violation parameter η . (c) Absorption images of the ^{87}Sr and ^{88}Sr atomic samples with and without resonant modulation.

can be interpreted as the amplitude of a finite-range mass-spin interaction [24], as a quantum-gravity property of the matter wave field [35], or as a gravitational mass tensor with a spin-dependent component in the standard model extension [36]. The Bloch frequency corresponds to the site-to-site energy difference induced by the gravitational force, and, according to the EP, the frequency difference $\delta_{87,88}$ for the two isotopes must depend only on the atomic mass ratio $R_{88,87} = m_{88}/m_{87}$ which is known with a relative uncertainty of 1.5×10^{-10} [37].

The experimental setup was based on an ultrahigh vacuum chamber in which the two Sr isotopes were alternately laser cooled and trapped [12]. An oven produced a thermal atomic beam which was slowed in a Zeeman slower and trapped in a magneto-optical trap (MOT) operating on the $^1S_0 - ^1P_1$ resonance transition at 461 nm. The loading time of the MOT was about 3 s and 7 s for ^{88}Sr and ^{87}Sr atoms, respectively. The temperature was further reduced by a second cooling stage in a “red” MOT operating on the $^1S_0 - ^3P_1$ intercombination transition at 689 nm. In the case of ^{87}Sr atoms, the cooling radiation (cycling on the $F = 9/2 \rightarrow F' = 11/2$ transition) was combined to a second “stirring” laser radiation (tuned on the $F = 9/2 \rightarrow F' = 9/2$ transition) to randomize the population of Zeeman sublevels to increase the trapping efficiency [38]. The red MOT confined about 5×10^6 ^{88}Sr atoms and 5×10^5 ^{87}Sr atoms with temperatures of 1 μK and 1.4 μK , respectively. The atoms were adiabatically loaded in a vertical optical lattice in 300 μs . For ^{87}Sr , this produced an unpolarized sample. The lattice potential was

generated by a single-mode frequency-doubled Nd:YVO₄ laser ($\lambda_L = 532$ nm) delivering up to 1.6 W on the atoms with two counterpropagating beams with a beam waist of about 300 μm . During the gravity measurements the lattice laser frequency was locked to a hyperfine component of molecular iodine by feedback to a piezomounted cavity mirror. The single-mode operation of the laser was monitored using a Fabry-Perot cavity; a self-referenced Ti:sapphire optical frequency comb enabled precise calibration of the laser frequency. The lattice depth U_0 was controlled and modulated by two acousto-optical modulators. The atomic cloud was imaged *in situ* at the end of each experiment cycle using resonant absorption imaging on a CCD camera with a spatial resolution of 5 μm .

We measured the Bloch frequency of ^{88}Sr and ^{87}Sr by applying an AM burst to the lattice depth for $t_M = 12$ s and 8 s at the $\ell = 2$ and $\ell = 1$ harmonic of ν_B , respectively, and thereafter detecting the resonant broadening of the atomic cloud width σ_z . A first set of measurements consisted of sweeping the AM frequency f_M to record a full resonance spectrum. The recording time for a whole resonance spectrum was about 1 h and led to a maximum resolution of 5×10^{-7} for $\nu_{B,88}$ and 1.6×10^{-6} for $\nu_{B,87}$. A typical resonant tunneling spectrum with the corresponding best fit is shown in Fig. 2. The error on the Bloch frequency determination was calculated as the standard error of the fit for each resonance profile.

In this work, we also demonstrated a new method to improve the precision of the measurement of ν_B and consequently of gravity acceleration by locking the AM oscillator frequency f_M to the Bloch frequency. In analogy to what is done in atomic clocks, f_M can be kept at the top

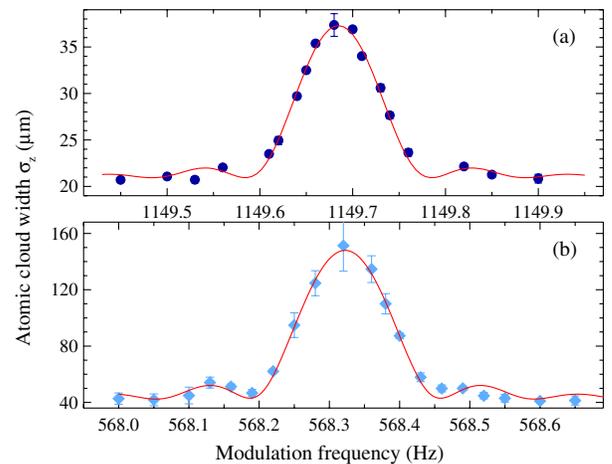


FIG. 2 (color online). Typical amplitude modulation spectra and the corresponding least-squares best-fit function (solid line) for (a) ^{88}Sr ($t_M = 12$ s, $\ell = 2$) and (b) ^{87}Sr ($t_M = 8$ s, $\ell = 1$, $\langle m_F \rangle = 0$) atoms. Both the lattice frequency and the lattice beam intensities were kept constant for each pair of measurements, while the modulation depth α was tuned to maximize the signal-to-noise ratio of each spectrum.

of the resonance spectrum in Fig. 2 by means of two consecutive AM interrogation cycles at each side of the spectrum. Subsequent demodulation was achieved by computing the difference of the two consecutive measurements of σ_z , which yielded an odd-symmetry error signal suitable for locking. The slope of the error signal across the resonance was about 0.6 mHz/ μm for typical experimental parameters ($\ell = 2$, $t_M = 10$ s and the modulation depth $\alpha = 6\%$ for ^{88}Sr , $\ell = 1$, $t_M = 6.8$ s and $\alpha = 4\%$ for ^{87}Sr). The Bloch frequency was determined by recording the f_M time series for about 700 s and taking the mean value of the time series. The sensitivity of the Bloch frequency measurement with the new method was characterized by its Allan deviation. Figure 3 shows the Allan deviation of a set of 101 recorded values of f_M for ^{88}Sr and a set of 42 values for ^{87}Sr . In both cases the Allan deviation scales as $t^{-1/2}$ (where t is the measurement time) with sensitivities at 1 s of $\sigma_{\nu_{B,88}} = 1.5 \times 10^{-6} \nu_{B,88}$ and $\sigma_{\nu_{B,87}} = 9.8 \times 10^{-6} \nu_{B,87}$, respectively. This new method allowed us to improve by more than 1 order of magnitude the sensitivity in the determination of the frequency of Bloch oscillations (and for gravity acceleration) for ^{88}Sr with respect to our previous results [12], achieving a precision of 5×10^{-8} for a single acquisition, while for ^{87}Sr we obtained a precision of 4×10^{-7} . The difference in precision between the two isotopes for both of the measurement techniques is due to the reduced signal-to-noise ratio in the absorption profile for ^{87}Sr . It is caused by the smaller natural abundance of this isotope and the presence of the 10-level hyperfine manifold that yields a higher Doppler temperature and a smaller (about a factor of 2) absorption cross section due to optical pumping in the imaging process and, for the frequency lock technique, a slightly higher cycle time (29 s versus 27 s).

Each pair of Bloch frequency measurements was used to determine the Eötvös ratio [39] given by

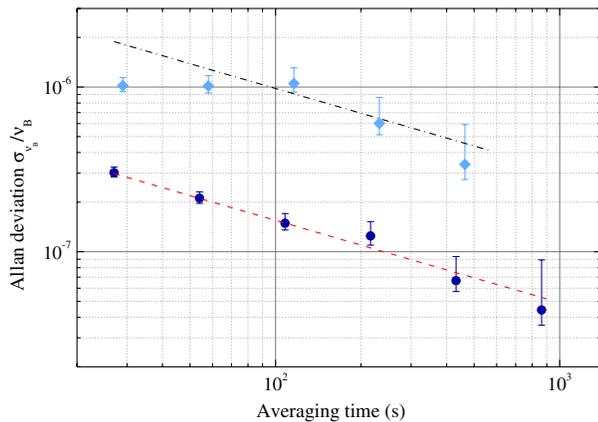


FIG. 3 (color online). Allan deviations of the Bloch frequency measurements for ^{88}Sr (circles) and ^{87}Sr (diamonds) and their corresponding $t^{-1/2}$ asymptotic behavior (lines) obtained by frequency locking the AM frequency generator to the coherent delocalization resonance, as described in the text.

$$\eta \equiv 2 \frac{a_{88} - a_{87}}{a_{88} + a_{87}} = 2 \frac{\nu_{B,88} - R_{88,87} \nu_{B,87}}{\nu_{B,88} + R_{88,87} \nu_{B,87}}, \quad (2)$$

where $a_i = 2h\nu_{B,i}/m_i\lambda_L$ ($i = 87, 88$) are the measured vertical accelerations for the two isotopes. The data were recorded in $N = 68$ measurement sessions. Figure 4(a) shows the experimental results for η , their average value, and the comparison with the null value predicted by general relativity. Each point η_i is determined with its own error σ_i given by the quadratic sum of the statistical error and the uncertainty on the systematic effects.

In our differential measurement, many systematic errors, such as misalignment of the lattice beams, lattice frequency calibration, gravity gradients, and Gouy phase shift, largely cancel and can be neglected at the present level of accuracy. The main contribution to the systematic error in local gravity measurement with trapped neutral atoms arises from the space-dependent lattice depth $U_0(z)$ due to the local intensity gradient of the two interfering Gaussian beams [33]. Since we are interested only in the effect of the gravity acceleration upon ν_B , the differential acceleration due to the residual intensity gradient must be removed from the ratio given in Eq. (2). The correction has been calculated to be

$$\Delta\eta_U = \frac{R_{88,87} - 1}{(\nu_{B,88} + R_{88,87}\nu_{B,87})/2} \frac{\partial_z U_0}{2\hbar k_0}, \quad (3)$$

where $k_0 = 2\pi/\lambda_L$ is the lattice laser wave number and we assumed that the difference in the trapping potential due to the dynamic polarizability of the two isotopes is negligible [40], so that $\partial_z U_0 = \partial_z U_{0,88} = \partial_z U_{0,87}$. The expression of the correction in Eq. (3) is then given by the product of the

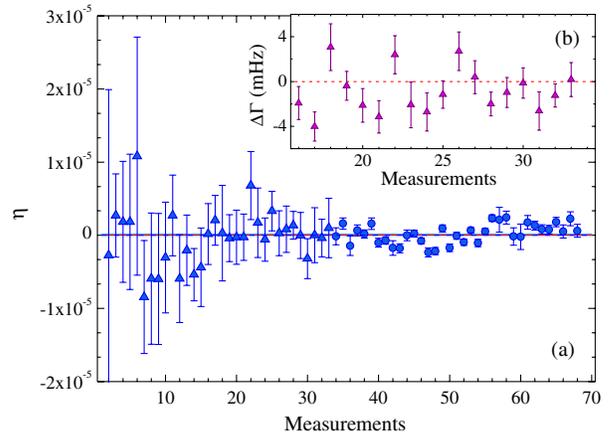


FIG. 4 (color online). Summary of the measurements for ^{87}Sr and ^{88}Sr Bloch frequency. (a) Measurements of the η parameter by AM resonant tunneling spectra (triangles) and by AM frequency lock (circles). The final weighted mean (blue dashed line) is compared with the null value expected from EP (red line). (b) Measurements of the resonance linewidth broadening $\Delta\Gamma$ for ^{87}Sr atoms. The dashed red line is $\Delta\Gamma = 0$.

shift of $\nu_{B,88}$ induced by the lattice beam gradient $\Delta\nu_U = \partial_z U_0/2\hbar k_0$ and a weight factor $R_{88,87} - 1 \sim 10^{-2}$ divided by the mean Bloch frequency $(\nu_{B,88} + R_{88,87}\nu_{B,87})/2$. The physical interpretation of Eq. (3) is that the acceleration due to the two-photon scattering process producing the confinement in the optical lattice has a reduced effect on the differential measurement but does not cancel out. This technical effect affects any EP tests employing an optical lattice [41]. A precise calibration of the acceleration due to the intensity gradient was done by measuring $\nu_{B,88}$ by means of the frequency lock technique. Repeated measurements of $\nu_{B,88}$ were performed with stabilized lattice frequency as a function of the total lattice power $P = P_1 + P_2 + 2\varepsilon\sqrt{P_1P_2}$, where P_1 and P_2 is the power sent to the atoms per beam and ε is a geometrical correction factor due to the mismatch of the width of the two beams of order unity. The resulting Bloch frequency shift was $\Delta\nu_U = (\partial\nu_B/\partial P)P = (6.16 \pm 0.56) \times 10^{-6} P$ Hz/mW, corresponding to $\Delta\eta_U \sim 3.6 \times 10^{-7}$ for typical operating conditions. The effect of magnetic field gradients in the differential ν_B measurement was carefully studied. Residual magnetic field gradients $b = \partial B/\partial z$ were estimated by a precise calibration of the ^{88}Sr red MOT vertical position dynamics to be less than $140 \mu\text{T} \cdot \text{m}^{-1}$. While ^{88}Sr is insensitive to magnetic field gradients at this level of precision [12], the sensitivity of the ^{87}Sr atomic sample depends on the average spin projection $\langle m_F \rangle$. It was estimated by applying a magnetic field gradient up to $210 \text{ mT} \cdot \text{m}^{-1}$ and measuring $\nu_{B,87}$, which resulted in a sensitivity factor $\partial\nu_{B,87}/\partial b = (2 \pm 15) \text{ mHz}/(\text{T} \cdot \text{m}^{-1})$, consistent with a null effect. The effect of tides was estimated to be less than 1×10^{-8} for a typical time interval of 1 h between the two ν_B measurements for the two isotopes. The total systematic uncertainty is thus dominated by the intensity gradient uncertainty at the level of 3×10^{-8} , while a residual lattice frequency error due to the frequency lock precision has been estimated to be lower than 1×10^{-8} .

The final result for the η parameter is $\eta = (0.2 \pm 1.6) \times 10^{-7}$, where the uncertainty corresponds to the standard deviation of the weighted mean $\sigma_{\bar{\eta}} = \sqrt{1/\sum_N(\sigma_i^{-2})}$, corrected by the reduced chi-square $[\chi^2/(N-1) = 2.78]$. In the case of unpolarized ^{87}Sr atoms, the mean contribution of kS_z is zero and $\eta = \beta_{88} - \beta_{87}$. This result can be interpreted in terms of the EP violation parameters for the fundamental constituents of the two atoms, according to different parametrizations [17,42], and it sets a 10^{-7} direct bound on the boson-to-fermion gravitational constant ratio f_{BF} from being different from 1 [29]. On the other hand, each ^{87}Sr spin component $S_z = I_z$ will feel different gravitational forces due to different spin-gravity coupling, as in the case of a magnetic field gradient, resulting in a broadening of the frequency response shown in Fig. 2. We analyzed a set of ^{87}Sr AM resonant tunneling spectra used for the

determination of η . The residual deviations of the measured linewidth Γ from the Fourier linewidth, after removing systematic broadening mechanisms such as the ones due to the two-body collisions and the residual magnetic field gradients, are shown in Fig. 4(b). The measured residual broadening $\Delta\Gamma = 0.4 \pm 0.5(\text{stat}) \pm 0.8(\text{syst})$ mHz sets an upper limit on the spin-gravity coupling $\Delta\Gamma = 2I_{87}k\ell\nu_{B,87}$, resulting in a spin-gravity coupling strength

$$k = (0.5 \pm 1.1) \times 10^{-7}.$$

Since the nucleus of ^{87}Sr has nine valence neutrons, this result also sets a limit of 10^{-7} for anomalous acceleration and spin-gravity coupling for the neutron either as a difference in the gravitational mass depending on the spin direction, which was previously limited at 10^{-23} [26], or as a coupling to a finite-range Leitner-Okubo-Hari Dass interaction, which was limited to less than 10 at 30 mm [24].

In conclusion, we performed a quantum test of EP for the bosonic ^{88}Sr isotope which has no spin versus the fermionic ^{87}Sr isotope which has a half-integer spin by coherent control of the atomic motion in an optical lattice under the effect of gravity. We obtained upper limits of $\sim 10^{-7}$ for pure inertial effects and for a possible spin-gravity coupling. The present results can set bounds for previously unmeasured parameters of the standard model extension [34,36]. Further enhancements in sensitivity will require the development of higher transferred momentum atom interferometry schemes for Sr atoms and simultaneous probing of the two isotopes [43]. Short-distance measurements ($r \leq 1$ cm) with $10^{-8}\nu_B$ precision can lower the limit of monopole-dipole interaction constants $g_p g_s$ of 9 orders of magnitude [27]. At the same time, Sr optical clocks are showing impressive advances in stability and accuracy with the possibility of building compact and transportable systems [44]. Results from a network of Sr optical clocks already set a limit to the coupling of fundamental constants to gravity [45]. It is possible then to envisage a future experiment in space where a Sr interferometer and a Sr optical clock would be operated at their limit performances to realize stringent tests of general relativity [46].

This project has received funding from INFN (MAGIA experiment) and from the EU's 7th FP under Grant Agreement No. 250072. We thank S. Capozziello, J. Tasson, and T. Zelevinsky for a critical reading of the manuscript and for useful suggestions.

*Present address: Department of Physics, Columbia University, 538 West 120th Street, New York, New York 10027-5255, USA.

†Also at: International Center for Theoretical Physics (ICTP), Trieste, Italy.

‡Guglielmo.Tino@fi.infn.it

- [1] C. W. Misner, K. S. Thorne, and J. A. Wheeler, *Gravitation* (Freeman, San Francisco, 1973).
- [2] D. Colladay and V. A. Kostelecký, *Phys. Rev. D* **55**, 6760 (1997).
- [3] T. Damour, F. Piazza, and G. Veneziano, *Phys. Rev. Lett.* **89**, 081601 (2002), and references therein.
- [4] S. Capozziello and M. De Laurentis, *Phys. Rep.* **509**, 167 (2011).
- [5] C. M. Will, *Living Rev. Relativity* **9**, 3 (2006).
- [6] A. D. Cronin, J. Schmiedmayer, and D. E. Pritchard, *Rev. Mod. Phys.* **81**, 1051 (2009).
- [7] *Atom Interferometry, Proceedings of the International School of Physics "Enrico Fermi," Course CLXXXVIII*, edited by G. M. Tino and M. A. Kasevich (Società Italiana di Fisica and IOS Press, Amsterdam, 2014).
- [8] S. Fray, C. A. Diez, T. W. Hänsch, and M. Weitz, *Phys. Rev. Lett.* **93**, 240404 (2004).
- [9] A. Bonnin, N. Zahzam, Y. Bidel, and A. Bresson, *Phys. Rev. A* **88**, 043615 (2013).
- [10] D. Schlippert, J. Hartwig, H. Albers, L. L. Richardson, C. Schubert, A. Roura, W. Schleich, W. Ertmer, and E. M. Rasel, *Phys. Rev. Lett.* **112**, 203002 (2014).
- [11] A. Peters, K. Y. Chung, and S. Chu, *Nature (London)* **400**, 849 (1999).
- [12] N. Poli, F.-Y. Wang, M. G. Tarallo, A. Alberti, M. Prevedelli, and G. M. Tino, *Phys. Rev. Lett.* **106**, 038501 (2011).
- [13] S. Dimopoulos, P. W. Graham, J. Hogan, and M. A. Kasevich, *Phys. Rev. Lett.* **98**, 111102 (2007).
- [14] G. M. Tino, F. Sorrentino, D. Aguilera, B. Battelier, A. Bertoldi, Q. Bodart, K. Bongs, P. Bouyer, C. Braxmaier, L. Cacciapuoti *et al.*, *Nucl. Phys. B, Proc. Suppl.* **243–244**, 203 (2013).
- [15] B. Altschul, Q. G. Bailey, L. Blanchet, K. Bongs, P. Bouyer, L. Cacciapuoti, S. Capozziello, N. Gaaloul, D. Giulini, J. Hartwig *et al.*, arXiv:1404.4307.
- [16] A. Kellerbauer, M. Amoretti, A. Belov, G. Bonomi, I. Boscolo, R. Brusa, M. Buchner, V. Byakov, L. Cabaret, C. Canali *et al.*, *Nucl. Instrum. Methods Phys. Res., Sect. B* **266**, 351 (2008).
- [17] M. A. Hohensee, H. Müller, and R. B. Wiringa, *Phys. Rev. Lett.* **111**, 151102 (2013); P. Hamilton, A. Zhmoginov, F. Robicheaux, J. Fajans, J. S. Wurtele, and H. Müller, *Phys. Rev. Lett.*, **112**, 121102 (2014).
- [18] F. W. Hehl, P. von der Heyde, G. D. Kerlick, and J. M. Nester, *Rev. Mod. Phys.* **48**, 393 (1976).
- [19] A. Peres, *Phys. Rev. D* **18**, 2739 (1978).
- [20] B. Mashhoon, *Classical Quantum Gravity* **17**, 2399 (2000).
- [21] Y. N. Obukhov, *Phys. Rev. Lett.* **86**, 192 (2001).
- [22] D. Bini, C. Cherubini, and B. Mashhoon, *Classical Quantum Gravity* **21**, 3893 (2004).
- [23] S. Capozziello, G. Lambiase, and C. Stornaiolo, *Ann. Phys. (Berlin)* **10**, 713 (2001).
- [24] W.-T. Ni, *Rep. Prog. Phys.* **73**, 056901 (2010), and references therein.
- [25] B. R. Heckel, E. G. Adelberger, C. E. Cramer, T. S. Cook, S. Schlamminger, and U. Schmidt, *Phys. Rev. D* **78**, 092006 (2008).
- [26] B. J. Venema, P. K. Majumder, S. K. Lamoreaux, B. R. Heckel, and E. N. Fortson, *Phys. Rev. Lett.* **68**, 135 (1992).
- [27] D. F. J. Kimball, I. Lacey, J. Valdez, J. Swiatlowski, C. Rios, R. Peregrina-Ramirez, C. Montcrieffe, J. Kremer, J. Dudley, and C. Sanchez, *Ann. Phys. (Berlin)* **525**, 514 (2013), and references therein.
- [28] D. J. Wineland, J. J. Bollinger, D. J. Heinzen, W. M. Itano, and M. G. Raizen, *Phys. Rev. Lett.* **67**, 1735 (1991).
- [29] J. D. Barrow and R. J. Scherrer, *Phys. Rev. D* **70**, 103515 (2004).
- [30] S. Herrmann, H. Dittus, and C. Lämmerzahl (QUANTUS and PRIMUS teams), *Classical Quantum Gravity* **29**, 184003 (2012).
- [31] G. Varoquaux, R. A. Nyman, R. Geiger, P. Cheinet, A. Landragin, and P. Bouyer, *New J. Phys.* **11**, 113010 (2009).
- [32] A. Alberti, G. Ferrari, V. V. Ivanov, M. L. Chiofalo, and G. M. Tino, *New J. Phys.* **12**, 065037 (2010).
- [33] M. G. Tarallo, A. Alberti, N. Poli, M. L. Chiofalo, F.-Y. Wang, and G. M. Tino, *Phys. Rev. A* **86**, 033615 (2012).
- [34] V. A. Kostelecký and J. D. Tasson, *Phys. Rev. D* **83**, 016013 (2011).
- [35] C. Lämmerzahl, *Classical Quantum Gravity* **15**, 13 (1998).
- [36] J. D. Tasson (private communication).
- [37] R. Rana, M. Höcker, and E. G. Myers, *Phys. Rev. A* **86**, 050502(R) (2012).
- [38] T. Mukaiyama, H. Katori, T. Ido, Y. Li, and M. Kuwata-Gonokami, *Phys. Rev. Lett.* **90**, 113002 (2003).
- [39] R. Eötvös, V. Pekár, and E. Fekete, *Ann. Phys. (Berlin)* **373**, 11 (1922).
- [40] T. Middelmann, S. Falke, C. Lisdat, and U. Sterr, *Phys. Rev. Lett.* **109**, 263004 (2012).
- [41] T. Kovachy, J. M. Hogan, D. M. S. Johnson, and M. A. Kasevich, *Phys. Rev. A* **82**, 013638 (2010).
- [42] T. Damour, *Classical Quantum Gravity* **13**, A33 (1996).
- [43] N. Poli, R. E. Drullinger, G. Ferrari, J. Léonard, F. Sorrentino, and G. M. Tino, *Phys. Rev. A* **71**, 061403(R) (2005).
- [44] N. Poli, C. W. Oates, P. Gill, and G. M. Tino, *Riv. Nuovo Cimento* **36**, 555 (2013).
- [45] S. Blatt, A. Ludlow, G. Campbell, J. Thomsen, T. Zelevinsky, M. Boyd, J. Ye, X. Baillard, M. Fouché, R. Le Targat *et al.*, *Phys. Rev. Lett.* **100**, 140801 (2008).
- [46] G. M. Tino, *Proceedings of the International School of Physics "Enrico Fermi," Course CLXXXVIII (Ref. [7])*.